

Cosmology with stacked cluster weak lensing and cluster-galaxy cross-correlations

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DOE-CSGF Program Review



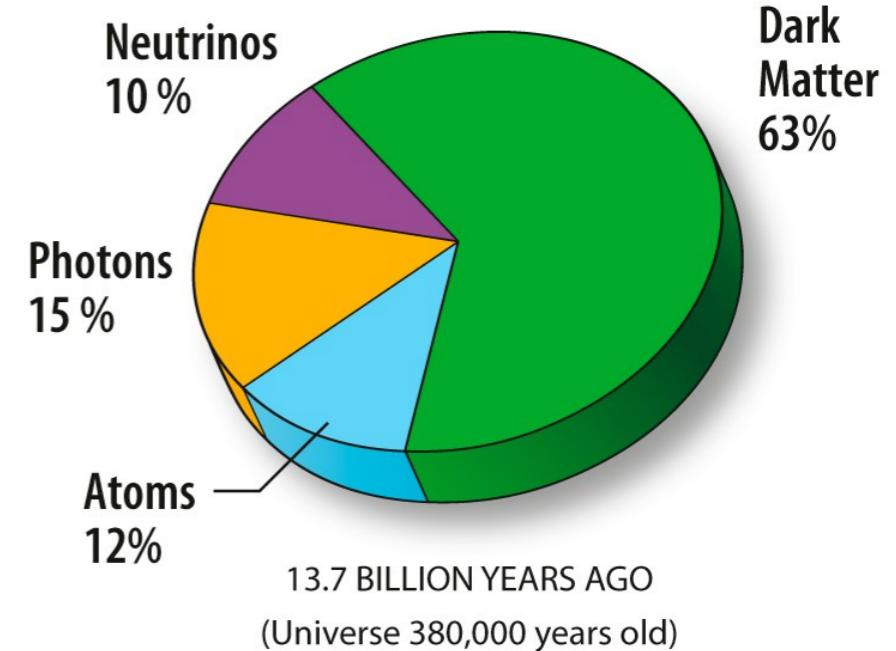
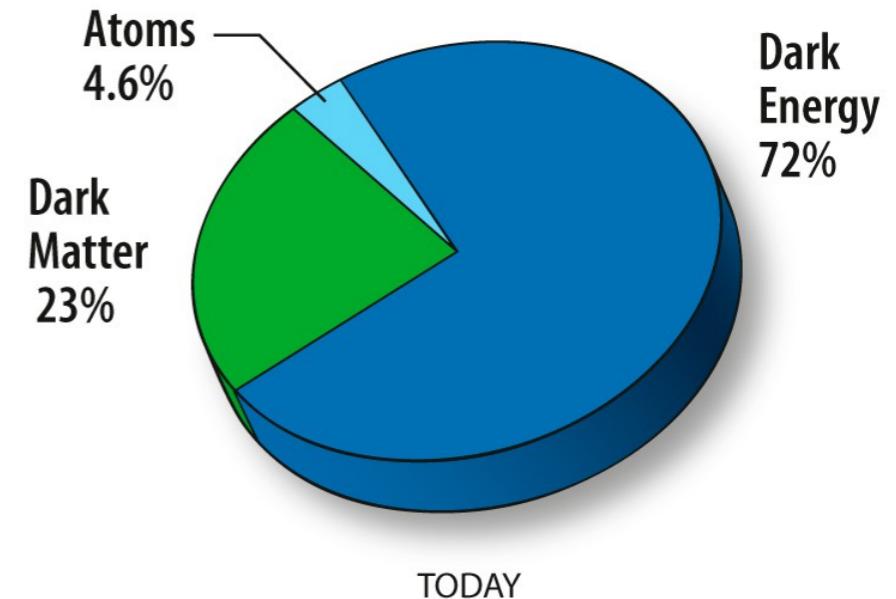
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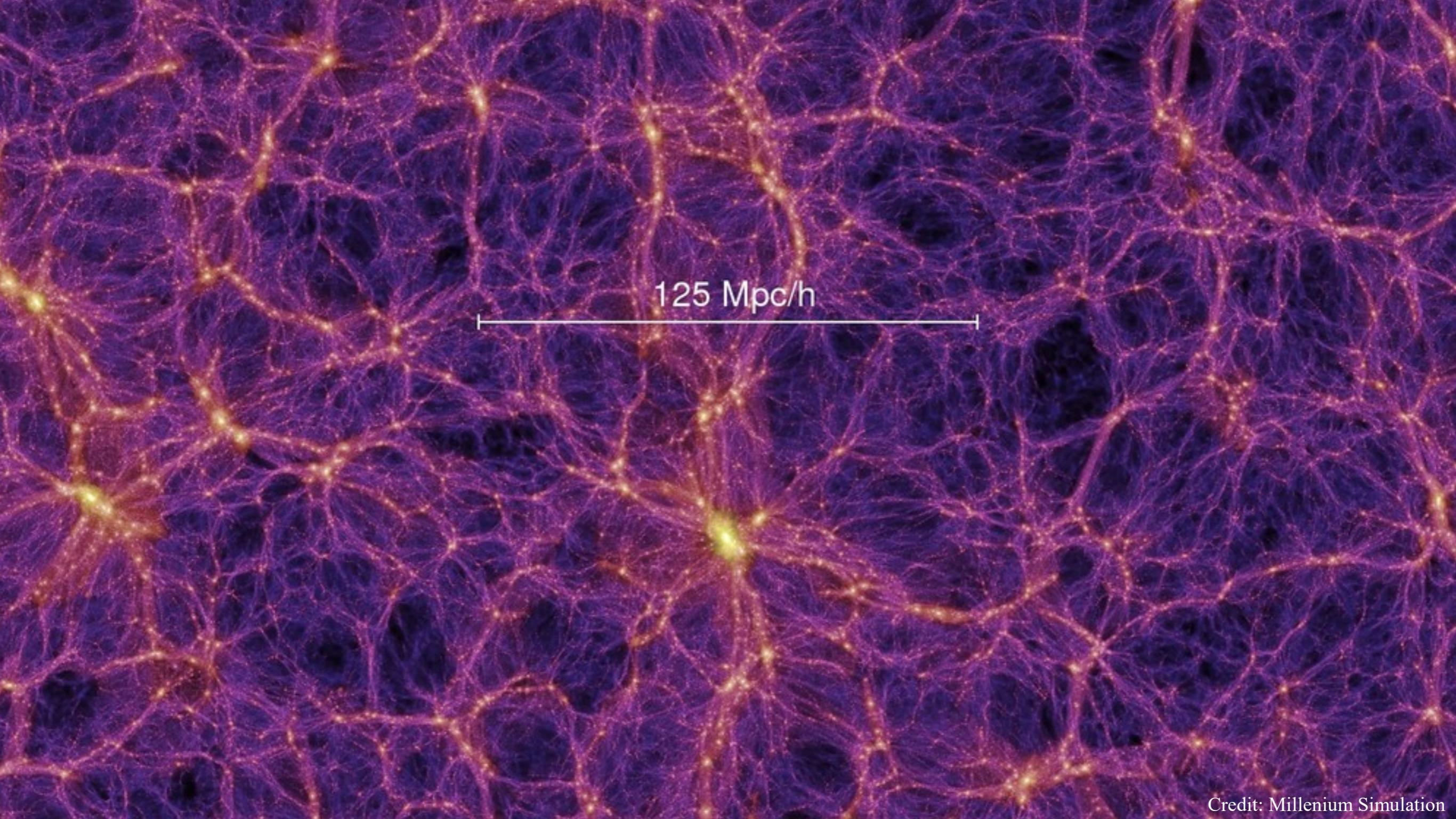


Physical cosmology

We live in a flat expanding universe with three major constituents:

1. **Baryonic matter**, what we are made of.
2. **Dark matter**, only interacts gravitationally.
3. **Dark energy**, an unknown form of energy that affects the expansion of the universe on large scales.





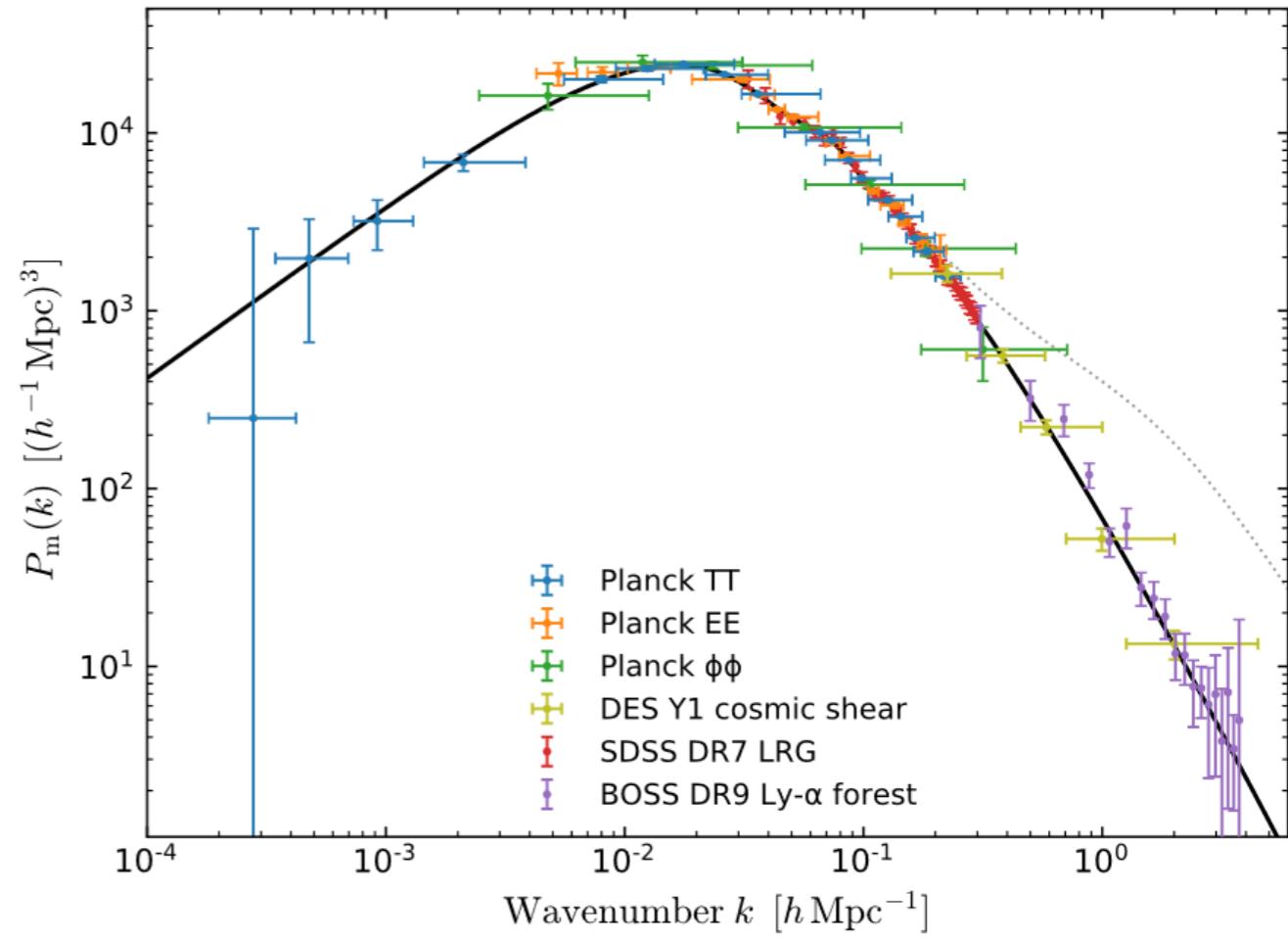
Credit: Millenium Simulation

Statistics of the cosmological density field

$$\Delta^2(k) = \frac{k^3 P(k)}{2\pi^2}$$

$$\sigma^2(R) = \int_0^\infty \Delta^2(k) |W_{\text{TH}}(k, R)|^2 d \ln k$$

$$\sigma_8 \equiv \sigma(R = 8 h^{-1} \text{ Mpc}, z = 0)$$

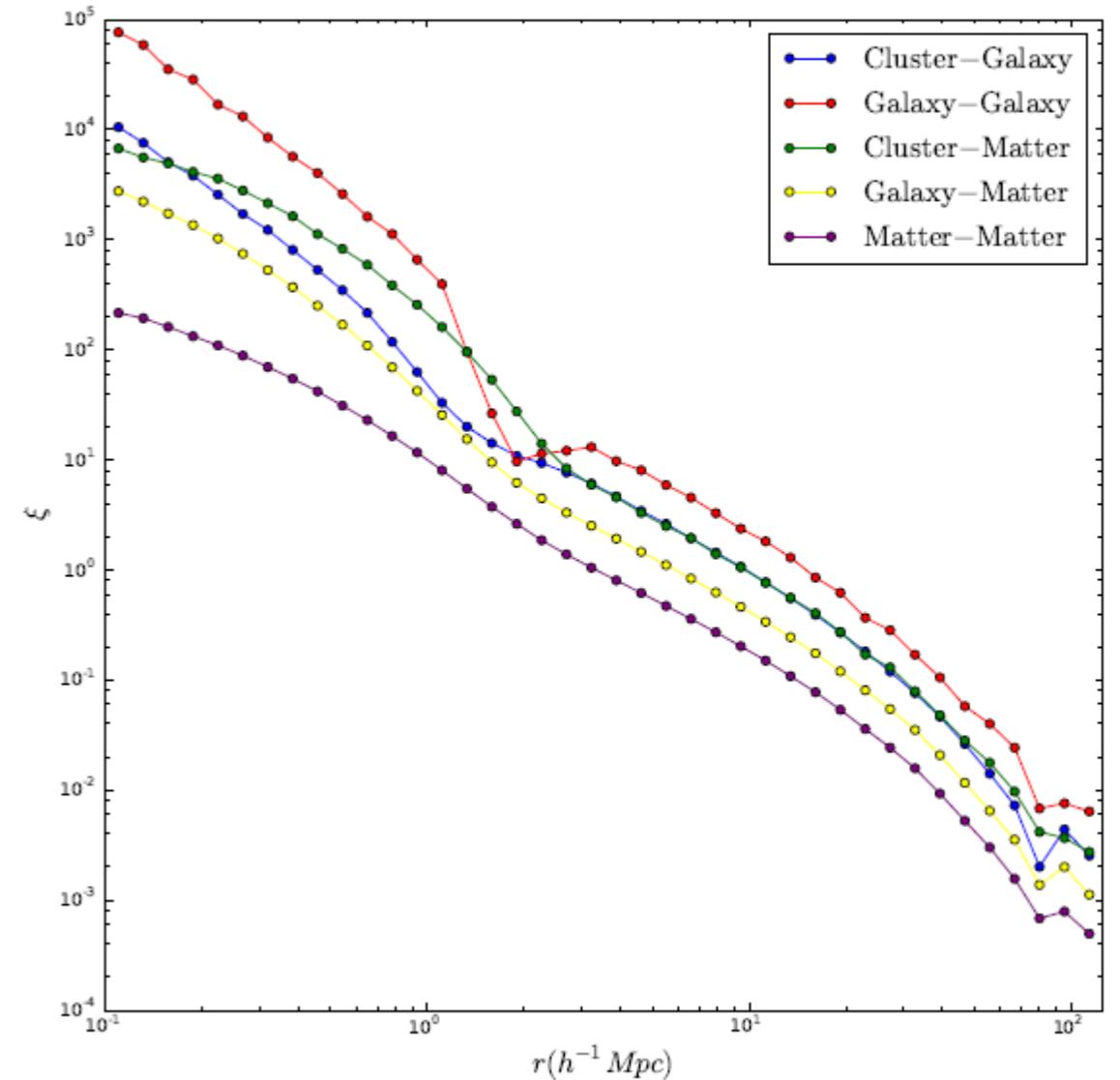


Statistics of the cosmological density field

$$\xi(r) = \frac{1}{2\pi^2} \int k^2 P(k) \frac{\sin(kr)}{kr} dk$$

$$\xi_{mm}(r) \propto \sigma_8^2$$

$$\xi_{AB}(r) = b_A(r)b_B(r)\xi_{mm}(r)$$

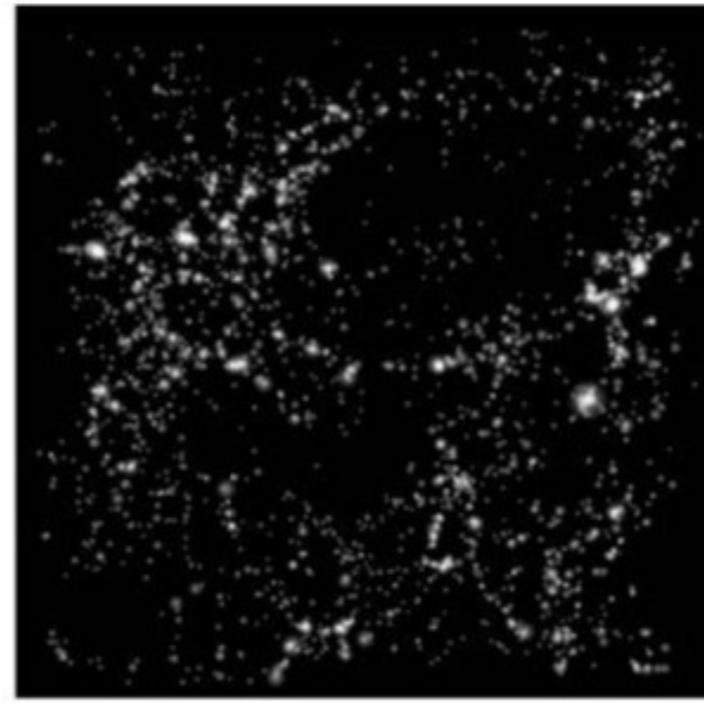


Biased tracers of an underlying distribution

Dark Matter

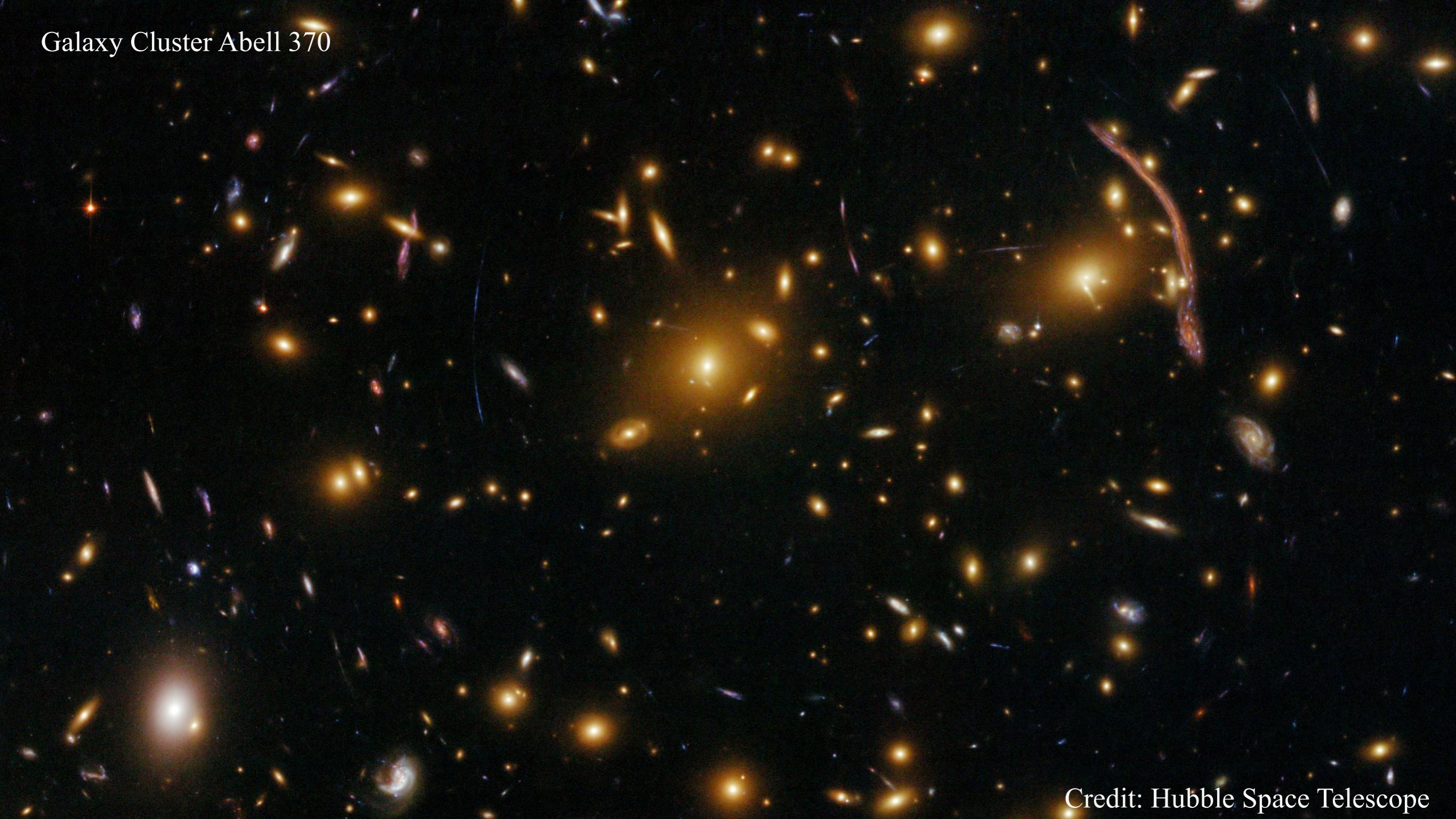


Galaxies



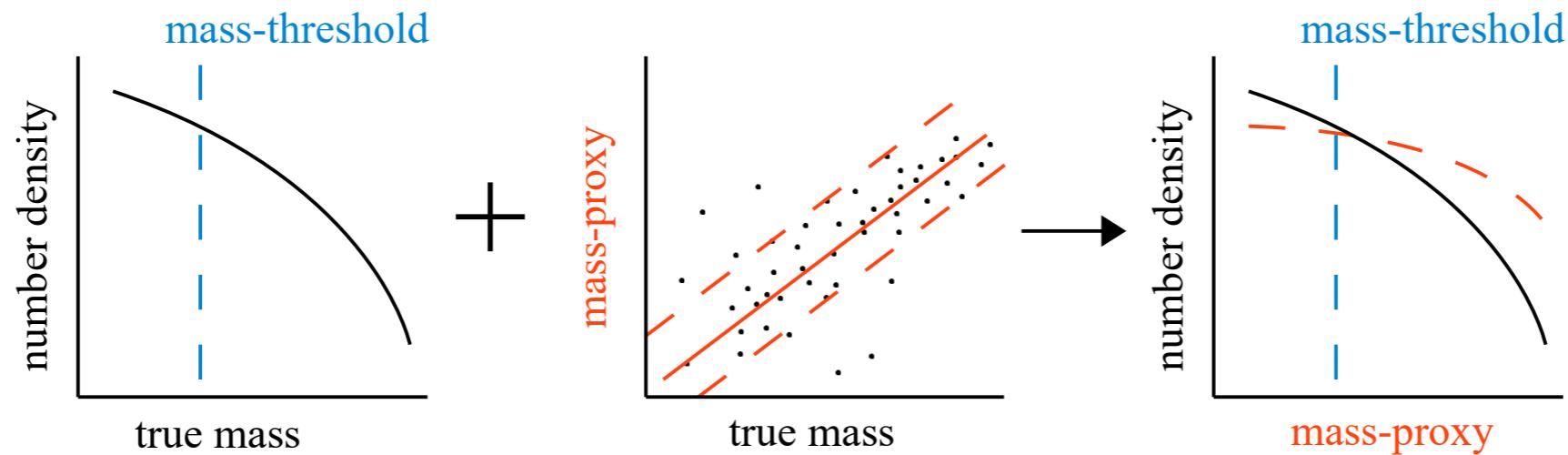
Visible galaxies trace the underlying distribution of invisible dark matter.

Galaxy Cluster Abell 370

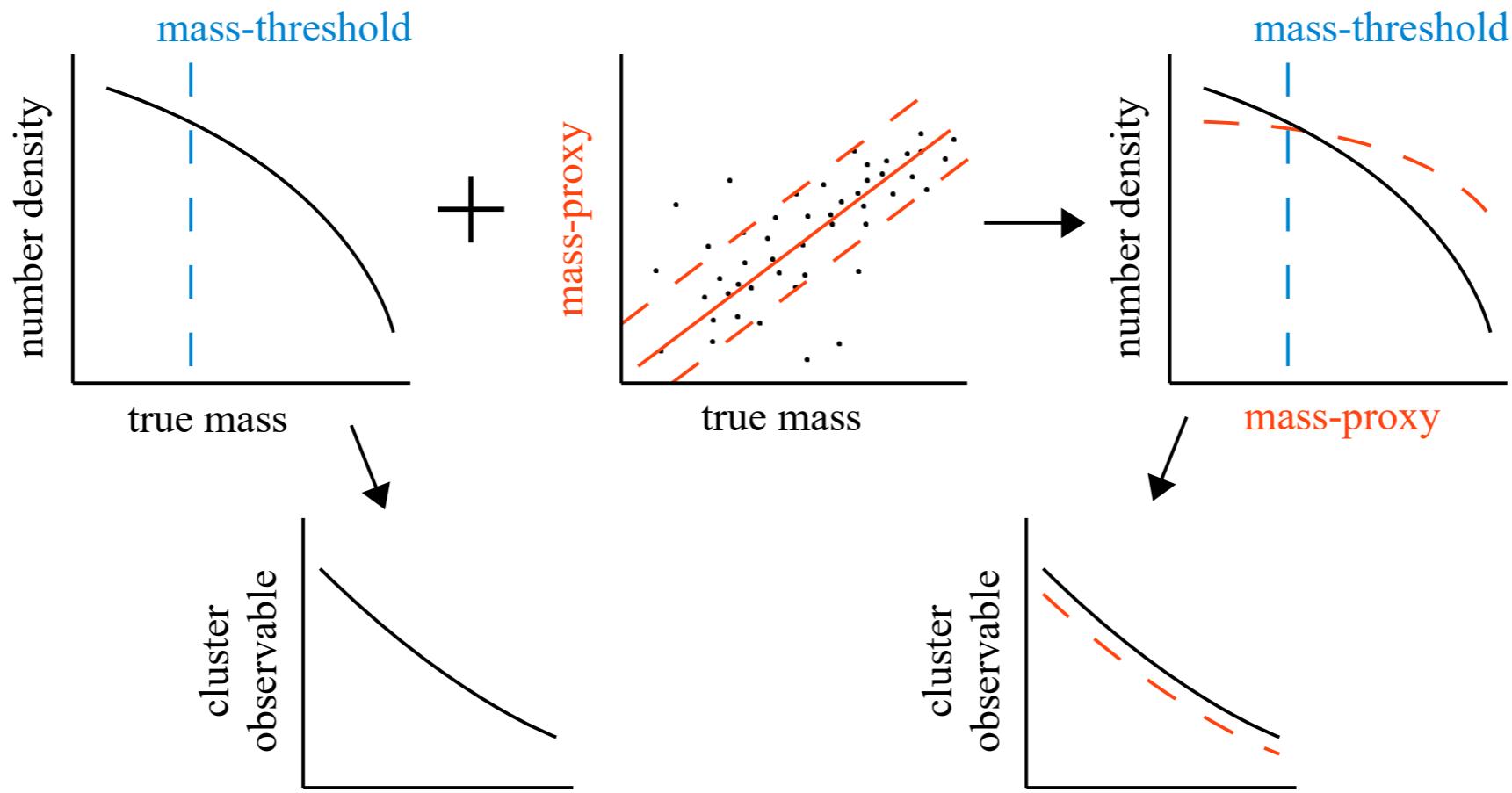


Credit: Hubble Space Telescope

Cluster mass-observable relation



Cluster mass-observable relation



Cluster and cosmological parameters of interest

σ_8 : amount of dark matter structure

Ω_m : density of the universe

$\sigma_{\ln M_c}$:log-normal scatter between cluster mass and observable

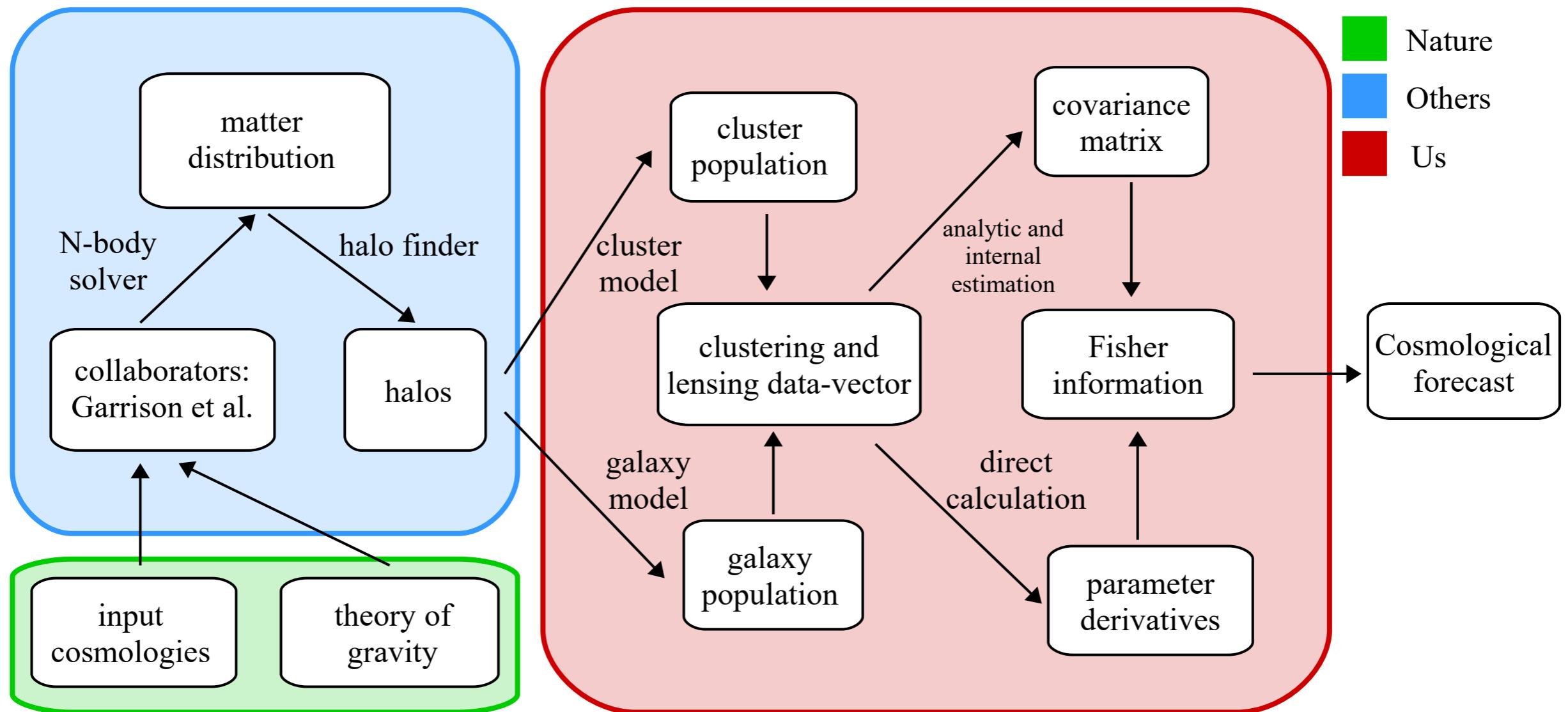
Three unknowns, three observables

Cluster weak lensing $\Delta\Sigma \propto \xi_{cm} \propto b_c \sigma_8^2,$

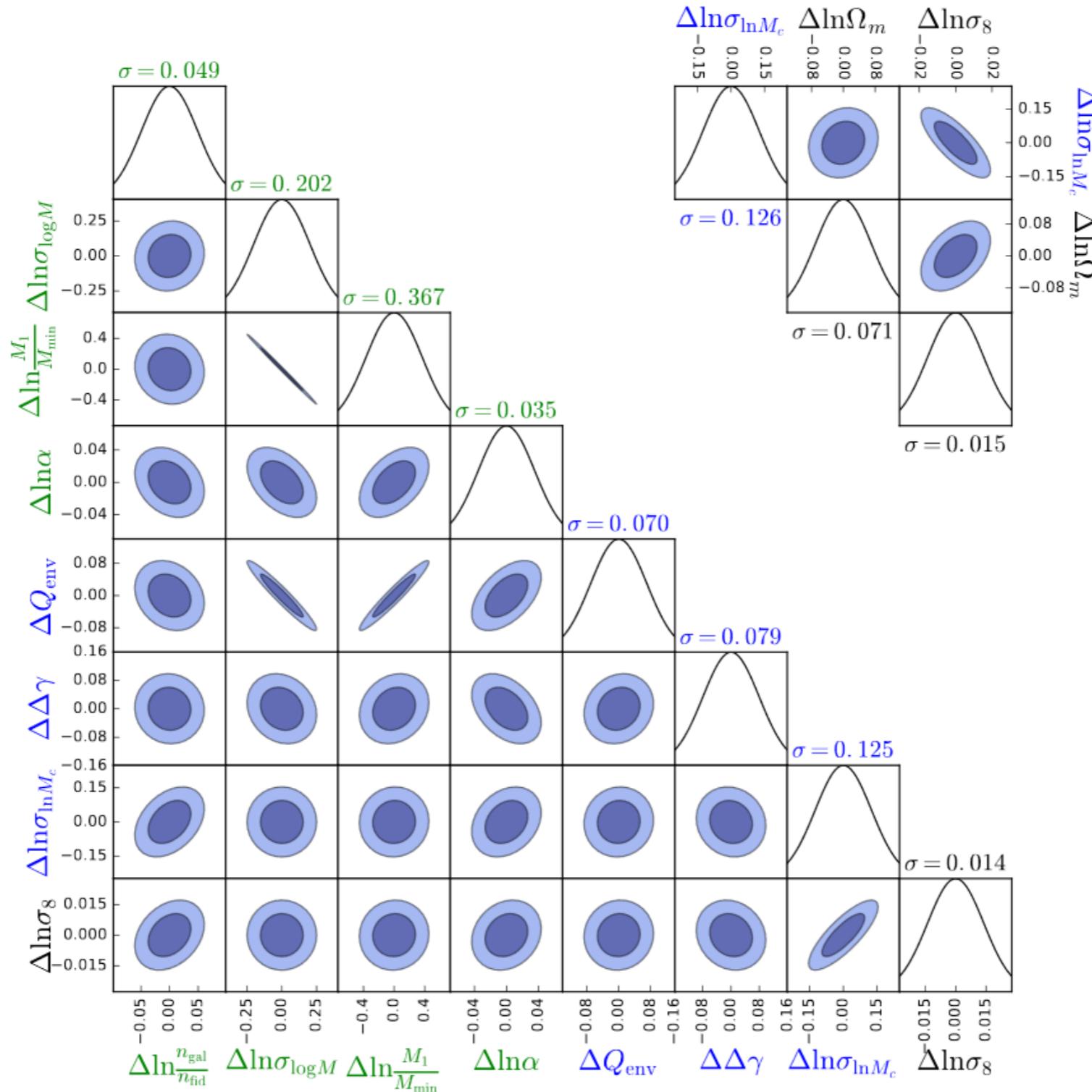
Cluster-galaxy
cross-correlation $w_{p,cg} \propto \xi_{cg} \propto b_c b_g \sigma_8^2,$

Galaxy auto-correlation $w_{p,gg} \propto \xi_{gg} \propto b_g^2 \sigma_8^2$

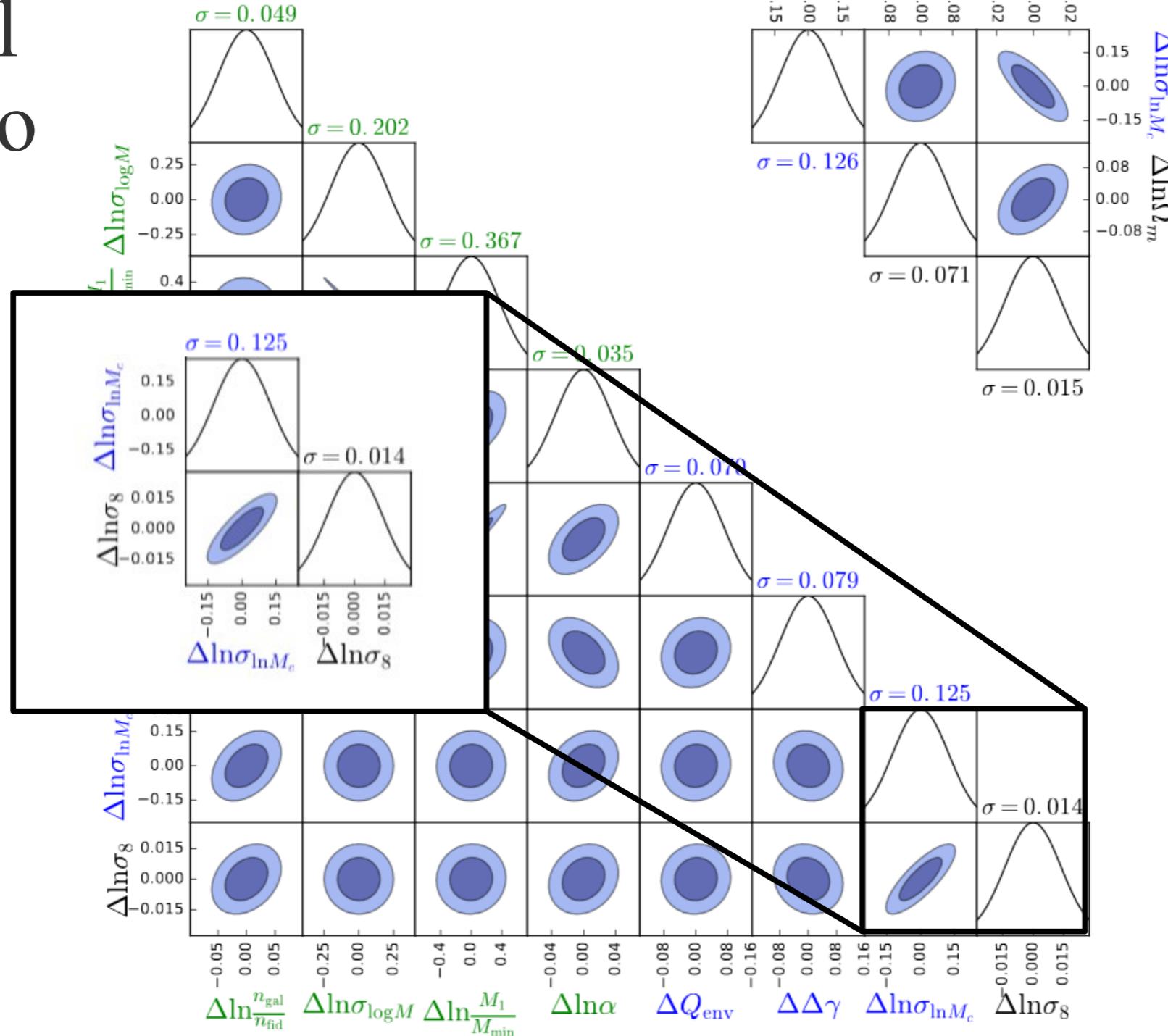
Analysis Overview



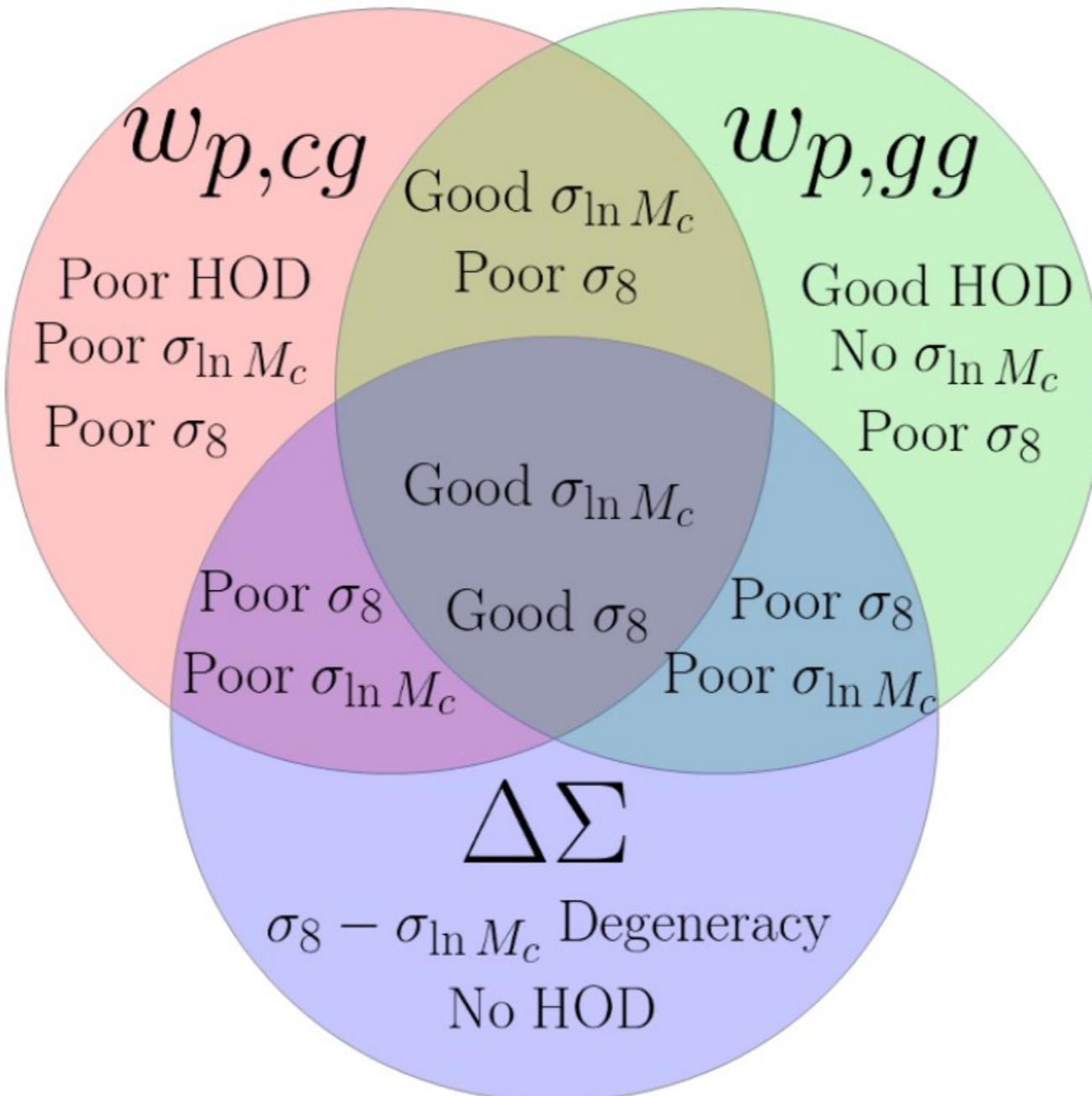
Fiducial Scenario



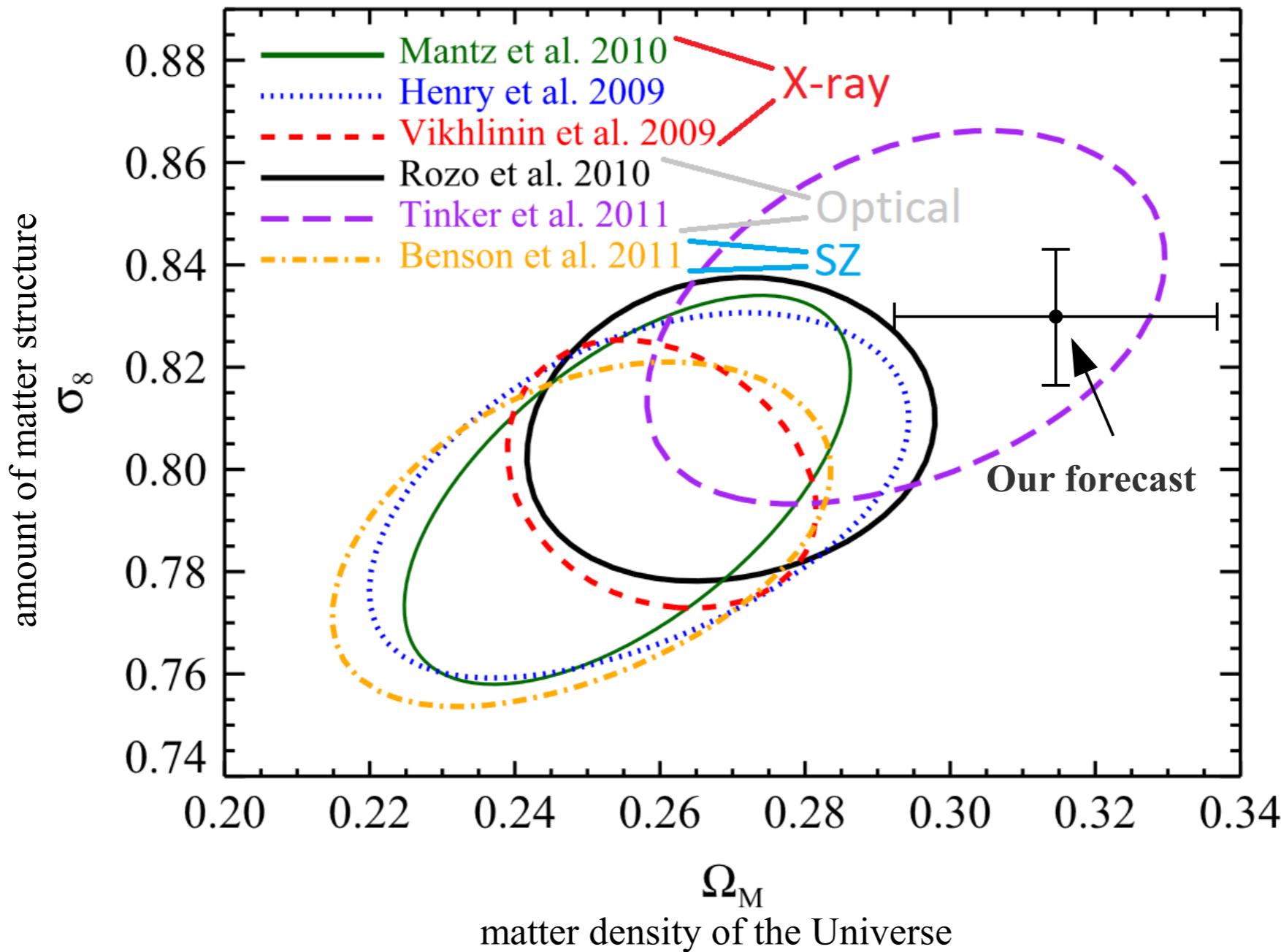
Fiducial Scenario



Where does the constraint come from?



Comparison with previous work





Acknowledgements and Questions